

Searching for Afterglow: Light Dark Matter Boosted by Supernova Neutrinos

Yen-Hsun Lin,^{4,1,*} Wen-Hua Wu,^{2,1} Meng-Ru Wu,^{1,3,4,†} and Henry Tsz-King Wong^{1,‡}

¹*Institute of Physics, Academia Sinica, Taipei 115, Taiwan*

²*Department of Physics, National Taiwan University, Taipei 106, Taiwan*

³*Institute of Astronomy and Astrophysics, Academia Sinica, Taipei 106, Taiwan*

⁴*Physics Division, National Center for Theoretical Sciences, Taipei 106, Taiwan*

A novel analysis is performed, incorporating *time-of-flight* (TOF) information to study the interactions of dark matter (DM) with standard model particles. After supernova (SN) explosions, DM with mass $m_\chi \lesssim \mathcal{O}(\text{MeV})$ in the halo can be boosted by SN neutrinos (SN ν) to relativistic speed. The SN ν boosted DM (BDM) arrives on Earth with TOF which depends only on m_χ and is independent of the cross section. These BDMs can interact with detector targets in low-background experiments and manifest as *afterglow* events after the arrival of SN ν . The characteristic TOF spectra of the BDM events can lead to large background suppression and unique determination of m_χ . New cross section constraints on $\sqrt{\sigma_{\chi e}\sigma_{\chi\nu}}$ are derived from SN1987a in the Large Magellanic Cloud with data from the Kamiokande and Super-Kamiokande experiments. Potential sensitivities for the next galactic SN with Hyper-Kamiokande are projected. This analysis extends the existing bounds on $\sqrt{\sigma_{\chi e}\sigma_{\chi\nu}}$ over a broad range of $r_\chi = \sigma_{\chi\nu}/\sigma_{\chi e}$. In particular, the improvement is by 1–3 orders of magnitude for $m_\chi < \mathcal{O}(100 \text{ keV})$ for $\sigma_{\chi e} \sim \sigma_{\chi\nu}$. Prospects of exploiting TOF information in other astrophysical systems to probe exotic physics with other DM candidates are discussed.

Introduction.—Although there is compelling evidence on the existence of dark matter (DM) as an additional gravity source, its properties and interactions remain unknown [1, 2]. Experimental searches of DM are intensely pursued worldwide [3–12]. Direct detection (DD) experiments focus on the weakly interacting massive particle (WIMP) scenario of DM mass $m_\chi \gtrsim \mathcal{O}(\text{GeV})$, with sensitivities approaching the neutrino floor [2]. The search for lighter WIMPs is an active area of research. One scenario with rapidly expanding interest is where light DM is up-scattered or boosted by known cosmic particles including baryons, electrons, and neutrinos [13–41]. The boosted DM (BDM) then carries kinetic energy T_χ much larger than when it is nonrelativistic (with velocity $v_\chi \sim 10^{-3}$) according to the Halo model. Nuclear and electron recoil events from BDM interaction with the detector targets will therefore have increased energy deposition, making DM with $m_\chi \lesssim \mathcal{O}(\text{GeV})$ experimentally accessible.

Time-of-flight (TOF) techniques are matured laboratory tools for differentiation or measurement of particle masses. This technique, however, has not been well exploited to probe exotic physics in astrophysical systems. One notable exception is the neutrino mass constraints derived from the timing distributions of supernova neutrinos (SN ν s) from SN1987a [42, 43]. We explore in this Letter a novel scenario of BDM with kinetic energy injected by SN ν interactions, and in particular where the prompt SN ν burst is also detected, providing a time-zero definition in terrestrial experiments. The prompt SN ν events will be followed by time-evolving BDM afterglow events where energy and time can be measured. The delay time between BDM and SN ν is a distinctive “smoking-gun” signature and provides unique information to infer m_χ , independent of the interaction cross section. Specifically, a delay time of $\Delta t \approx$

$10 \text{ days} \times [R/(8 \text{ kpc})][m_\chi/(10 \text{ keV})]^2 [T_\chi/(10 \text{ MeV})]^{-2}$ for SN ν BDM traveling an astronomical distance R before reaching the Earth highlights that although BDM has $v_\chi \sim c$, the delay can be substantial but measurable in a duration post the arrival of SN ν . In contrast, most proposed BDM scenarios rely on steady sources, e.g., cosmic rays [13–17, 19–29, 32, 35–37], stellar ν [18, 30], diffuse SN ν [31, 34], etc, for which the BDM flux is constant with time and lacks any time-dependent feature.

We explore the signatures of SN ν BDM with SN1987a in the Large Magellanic Cloud (LMC) and a future supernova (SN) in the Galactic Center (GC) to derive the fluxes and the associated electron-recoil event rates via $\sigma_{\chi e}$ in multikiloton water Cherenkov detectors, including Kamiokande, Super-Kamiokande (Super-K), and Hyper-Kamiokande. The scenario of SN ν BDM depends on finite DM cross section with ν ($\sigma_{\chi\nu}$), which may originate from an effective Lagrangian $\bar{\chi}\Gamma\chi\bar{\ell}_i\Gamma\ell_i/\Lambda$ where χ and $\ell_i = (\nu_i, i)$ are the DM and SM fields with $i = e, \mu, \tau$. The vertex Γ denotes the interaction type and Λ indicates certain cutoff scale. Possible interactions between χ with ν is a subject of intense recent interest [44–51]. They can naturally arise in many particle physics models such as the extensively studied $B - L$ and $L_\mu - L_\tau$, where the new gauge bosons can kinematically mix with the standard model photon. Further constraints will be provided by this work.

DM boosted by SN ν .—Assuming a SN explodes near the center of a galaxy (location O in Fig. 1), it emits a large amount of $\mathcal{O}(10)$ MeV neutrinos within $\tau \approx 10$ s carrying total luminosity $L_{\nu,\text{tot}} \approx 3 \times 10^{52} \text{ erg s}^{-1}$. We approximate these SN ν by an expanding thin spherical shell with a radius r away from O and a thickness $d \approx c\tau$ (see Fig. 1). The radially propagating SN ν within the

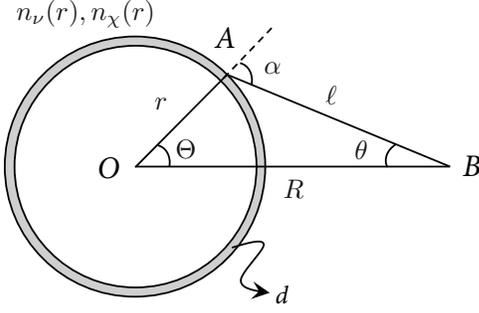


FIG. 1. Schematic plot of DM boosted by $\text{SN}\nu$ within an expanding spherical shell with width d at radius r . The SN occurs at O . BDM from A arrives B with a scattering angle α .

shell has a number density of

$$\frac{dn_\nu}{dE_\nu} = \sum_i \frac{L_{\nu_i}}{4\pi r^2 \langle E_{\nu_i} \rangle} E_\nu^2 f_{\nu_i}(E_\nu), \quad (1)$$

where $L_{\nu_i} = L_{\nu, \text{tot}}/6$ is the luminosity of each flavor (ν_e, ν_μ, ν_τ and their antineutrinos). We take the average energy $\langle E_{\nu_e} \rangle, \langle E_{\bar{\nu}_e} \rangle$, and $\langle E_{\nu_x} \rangle$ ($\nu_x \in \{\nu_\mu, \nu_\tau, \bar{\nu}_\mu, \bar{\nu}_\tau\}$) to be 11, 16, 25 MeV, respectively [52]. The energy distribution follows a Fermi-Dirac distribution f_{ν_i} with a pinch parameter $\eta_{\nu_i} \equiv \mu_{\nu_i}/T_{\nu_i} = 3$, such that $T_{\nu_i} \approx \langle E_{\nu_i} \rangle/3.99$.

With a nonvanishing DM- ν interaction, these neutrinos can upscatter DM in the halo [with number density $n_\chi(r)$] when they propagate outward. The BDM from location A can reach the Earth at B (with a distance R away from the center) with a scattering angle α after traveling a length ℓ . At neutrino energy E_ν much larger than the typical DM kinetic energy in the halo, DM can be approximated as at rest, and the BDM kinetic energy is given by

$$T_\chi = \frac{E_\nu^2}{E_\nu + m_\chi/2} \left(\frac{1 + \cos \theta_c}{2} \right), \quad (2)$$

where $\theta_c \in [0, \pi]$ is the scattering angle in the center-of-mass (c.m.) frame. One can relate θ_c to the lab frame scattering angle $\alpha \in [0, \pi/2]$ by $\theta_c = 2 \tan^{-1}(\gamma \tan \alpha)$ and $\gamma = (E_\nu + m_\chi)/\sqrt{m_\chi(2E_\nu + m_\chi)}$. Assuming $\sigma_{\chi\nu}$ is independent of θ_c in the c.m. frame, the normalized BDM angular distribution in the lab frame is given by

$$f_\chi(\alpha, E_\nu) = \frac{\gamma^2 \sec^3 \alpha}{\pi(1 + \gamma^2 \tan^2 \alpha)^2}, \quad (3)$$

such that $\int d\Omega_\alpha f_\chi(\alpha, E_\nu) = 1$ for any given E_ν , where $d\Omega_\alpha = 2\pi \sin \alpha d\alpha$. In Fig. 2, we plot $2\pi \sin \alpha f_\chi(\alpha)$ for a fixed $T_\chi = 10$ MeV (corresponding to different E_ν) with different m_χ . It shows that for BDM with $m_\chi/T_\chi \ll 1$, they are confined within a small scattering angle relative to the direction of $\text{SN}\nu$.

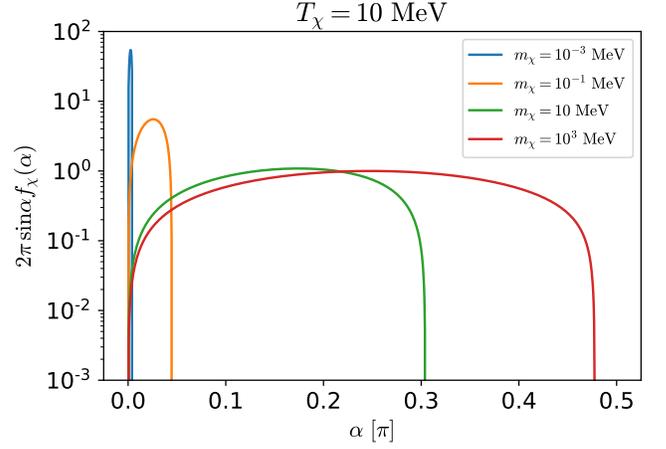


FIG. 2. The BDM angular distribution $f_\chi(\alpha)$ times $2\pi \sin \alpha$ for fixed $T_\chi = 10$ MeV and different m_χ .

The BDM emissivity j_χ at location A can be written as

$$j_\chi(r, T_\chi, \alpha) = c\sigma_{\chi\nu}n_\chi \left(\frac{dn_\nu}{dE_\nu} \right) \left(\frac{dE_\nu v_\chi}{dT_\chi c} \right) f_\chi, \quad (4)$$

where the BDM velocity $v_\chi/c = \sqrt{T_\chi(2m_\chi + T_\chi)}/(m_\chi + T_\chi)$, and can be evaluated using Eqs. (1) to (3).

Time-dependent BDM flux at Earth.—To obtain the BDM flux (number of BDM per unit time per unit energy per solid angle) at Earth $d\Phi_\chi/(dT_\chi d\Omega)$ (location B in Fig. 1), we shall integrate all j_χ along the line of sight ℓ ,

$$\frac{d\Phi_\chi(T_\chi, \theta, t')}{dT_\chi d\Omega} = \int d\ell j_\chi(r, T_\chi, \alpha) H \left(t' - \frac{r}{c} - \frac{\ell}{v_\chi} \right) \times H \left(\frac{r}{c} + \frac{\ell}{v_\chi} + \tau - t' \right), \quad (5)$$

where $d\Omega = 2\pi \sin \theta d\theta$ is viewed from B . The Heaviside functions limit j_χ to being nonzero only within the spherical shell of width d where $\text{SN}\nu$ are present. The arrival time of BDM, t' , relative to the time of SN explosion, includes the propagation time of $\text{SN}\nu$ from O to A (r/c) and the traveling time of BDM from A to B (ℓ/v_χ).

Integrating Eq. (5) over $d\Omega$ and approximating $H(x - x_0)H(x_0 + \epsilon - x) \sim \epsilon \delta(x_0)$ for $\epsilon \ll x_0$, we obtain

$$\begin{aligned} \frac{d\Phi_\chi(T_\chi, t')}{dT_\chi} &= 2\pi\tau \int d \cos \theta d\ell j_\chi(r, T_\chi, \alpha) \delta \left(t' - \frac{r}{c} - \frac{\ell}{v_\chi} \right) \\ &= 2\pi\tau \int_0^1 d \cos \theta \mathcal{J} j_\chi(r, T_\chi, \alpha) \Big|_{t' = \frac{r}{c} + \frac{\ell}{v_\chi}}, \end{aligned} \quad (6)$$

where

$$\mathcal{J} = \left(\frac{\ell - R \cos \theta}{rc} + \frac{1}{v_\chi} \right)^{-1} \quad (7)$$

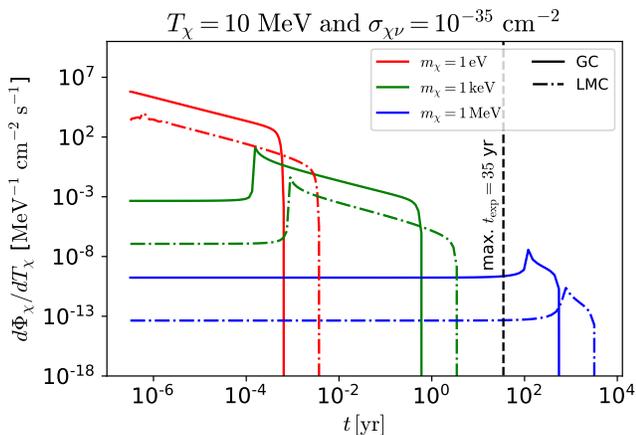


FIG. 3. The BDM flux at Earth vs. t with different m_χ for $T_\chi = 10$ MeV and $\sigma_{\chi\nu} = 10^{-35}$ cm 2 . Fluxes resulting from a SN in the GC and from SN1987a in LMC are shown with solid and dash-dotted lines. The black dashed line indicates the maximum exposure time $t_{\text{exp}} = 35$ years (see text for details).

appears due to the change of variable $d\ell = \mathcal{J}dt'$. Note that for a given (t', θ) , one can find a unique solution of (r, ℓ, α) and compute the integration.

BDM flux from SN in the GC and LMC.—We now compute the BDM fluxes at the Earth from SN1987a in LMC and from a SN in the GC. We characterize n_χ in the Milky Way (MW) and LMC by Navarro-Frenk-White (NFW) and Hernquist profiles respectively. Both share the same expression

$$n_\chi(r) = \frac{\rho_s}{m_\chi} \frac{1}{r_s \left(1 + \frac{r}{r_s}\right)^n}, \quad (8)$$

with $(n, \rho_s, r_s) = (2, 184 \text{ MeV cm}^{-3}, 24.4 \text{ kpc})$ for MW [53] and $(n, \rho_s, r_s) = (3, 68 \text{ MeV cm}^{-3}, 31.9 \text{ kpc})$ for LMC [54]. The distances R for the two are $(R_{\text{GC}}, R_{\text{LMC}}) = (8.5, 50) \text{ kpc}$. We neglect the contribution from $r < 10^{-5} \text{ kpc}$ since the profile in the inner region is highly uncertain and the adopted profile diverges when $r \rightarrow 0$.

Fig. 3 shows $d\Phi_\chi/dT_\chi$ versus $t > \tau$ for $T_\chi = 10$ MeV with different m_χ for SN in the GC (solid) and in LMC (dot-dashed), assuming $\sigma_{\chi\nu} = 10^{-35}$ cm 2 . Note that we define a shifted time coordinate $t = t' - R/c$ as the delayed arrival time for BDM relative to SN ν . For $m_\chi = 1$ keV and 1 MeV, the most prominent feature is that the BDM fluxes contain a rising part and peak at $t_p \approx R(1/v_\chi - 1/c)$. This is mainly due to the increase of $n_\chi \propto r^{-1}$ toward the halo center. The postpeak tails are due to BDM contributions with larger scattering angles. For $m_\chi = 1$ eV, $t_p \approx 0.004$ s is too short and overlaps with the 10 s duration of SN ν to be shown in Fig. 3. Comparing BDM fluxes coming from the GC to LMC, the LMC cases have smaller fluxes and larger t_p due to

larger R and smaller halo density.

Fig. 3 also shows another important feature—the BDM flux for a given T_χ and m_χ vanishes after some time post t_p , which is related to the sharp cutoff of f_χ shown in Fig. 2. This allows us to consider a reduced duration for BDM searches after the arrival of SN ν . Practically, a detector that can probe BDM has a threshold energy T_{th} , below which the detector is insensitive to BDM. Thus, for a given m_χ , one can define the latest possible arrival time of BDM with $T_\chi = T_{\text{th}}$ as the vanishing time t_{van} to analyze the data. We stress that all these time-dependent features only depend on m_χ but not $\sigma_{\chi\nu}$. Consequently, if such BDM is detected, analyzing the time profile of the signal will allow direct measurements of m_χ .

Events in Kamiokande and Super-K.—For BDM that also interact with electrons with a cross section $\sigma_{\chi e}$, they can produce signals in neutrino or DM experiments. The total event number N_χ induced by BDM with $T_{\text{th}} \leq T_\chi \leq T_{\text{max}}$ within an exposure time $t_0 \leq t \leq t_{\text{exp}}$ is given by

$$N_\chi = N_e \sigma_{\chi e} \int_{T_{\text{th}}}^{T_{\text{max}}} dT_\chi \int_{t_0}^{t_{\text{exp}}} dt \epsilon \frac{d\Phi_\chi}{dT_\chi}, \quad (9)$$

with N_e the total target number of electrons and ϵ the signal efficiency. We consider the water Cherenkov experiments, Kamiokande and Super-K, to calculate N_χ for BDM from LMC (by SN1987a) and from a SN in GC. They have $N_e = (M_T/m_{\text{H}_2\text{O}})N_A n_e$ with M_T the fiducial detector mass, $m_{\text{H}_2\text{O}}$ the water molar mass, N_A the Avogadro constant and n_e the electron number per water molecule. We take $M_T = 2.2$ and 22.5 kton for Kamiokande and Super-K, respectively [55, 56], and set $(T_{\text{th}}, T_{\text{max}}) = (5, 100) \text{ MeV}$ for both. We make a conservative choice of taking $\epsilon = 50\%$, lower than the energy-dependent efficiency roughly ranging from 50% to 75% reported for solar ν detection in Super-K [56]. For signal duration, we consider $t_0 = 10$ s to approximately exclude events produced by SN ν , and let $t_{\text{exp}} = \min(t_{\text{van}}, t_{\text{cut}} = 35 \text{ yrs})$ depending on m_χ . For the LMC case, the considered duration thus includes the running time of Kamiokande from 1987 to 1996 and Super-K after 1996 for heavier m_χ . For the GC case, we consider Super-K only. The main background for both comes from the solar and atmospheric neutrinos for $T_\chi \lesssim 20$ MeV and $T_\chi \gtrsim 20$ MeV, respectively. We adopt values in Table XIV of Ref. [56] and the FLUKA simulation result in Ref. [57] to estimate the background.

Fig. 4 shows N_χ vs. m_χ resulting from the GC and LMC given $\sigma_{\chi\nu} = \sigma_{\chi e} = 10^{-35}$ cm 2 . We first discuss the GC case where only Super-K is considered. The red-solid dots show that $N_\chi \propto m_\chi^{-1}$ perfectly for $m_\chi \leq 25$ keV, which corresponds to having $t_{\text{exp}} = t_{\text{van}} \leq t_{\text{cut}}$. This is because for smaller m_χ , all BDM arrive at the detector before t_{cut} so that N_χ is proportional to the amount of DM in the halo. For heavier m_χ , however, a larger part of BDM flux only arrives after t_{cut} (see Fig. 3), leading

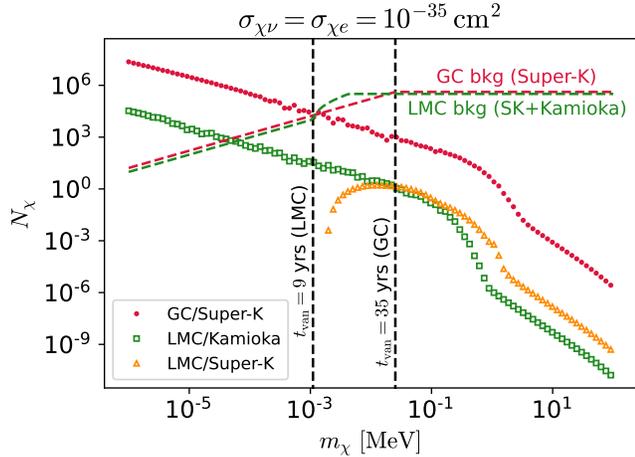


FIG. 4. BDM events in water Cherenkov detectors N_χ as a function of m_χ for both the GC and LMC (SN1987a) cases. For LMC, events in Kamiokande (hollow-green squares) and Super-K (hollow-orange triangles) are shown separately. For the GC, only Super-K (red dots) is considered. Background counts (dashed lines) are also shown for both cases.

to a faster decrease of N_χ with increasing m_χ . For the same reason, the background counts (red-dashed curve) $N_b \simeq 526 M_T t_{\text{exp}}$ for $m_\chi \leq 25$ keV due to a constant background rate of ~ 526 events per kton per year [56]. For $m_\chi > 25$ keV where $t_{\text{exp}} = t_{\text{cut}}$ is applied, N_b stays constant.

For BDM associated with SN1987a in LMC, we plot N_χ in Kamiokande (1987–1996) and Super-K (after 1996) by hollow green squares and orange triangles separately. The behavior of $N_\chi(m_\chi)$ in Kamiokande is similar to that of the GC case, but falls off faster for large m_χ due to the maximal exposure time of 9 yr only. The difference at small m_χ is mainly due to different detector fiducial mass M_T , geometric dilution factor $1/R^2$, and the characteristic density ρ_s of DM profiles. A simple estimate gives $N_\chi^{\text{GC}}/N_\chi^{\text{LMC}} \sim (\rho_s^{\text{NFW}}/\rho_s^{\text{Hernquist}})(R_{\text{LMC}}^2/R_{\text{GC}}^2)(M_T^{\text{SK}}/M_T^{\text{Kamioka}}) \sim \mathcal{O}(10^3)$ consistent with Fig. 4. Super-K here only starts to accumulate events for $m_\chi \gtrsim 1.1$ keV whose $t_{\text{van}} > 9$ yr, and eventually dominates the contribution to N_χ more than that from Kamiokande for larger m_χ . For comparison, we also plot the combined background numbers N_b from both detectors.

Sensitivity and constraint.—We use N_χ and N_b derived above to estimate the constraint and sensitivity on light DM, taking for simplicity

$$n_\sigma = \frac{N_\chi}{\sqrt{N_\chi + N_b}}. \quad (10)$$

The dependencies of sensitivity (s) versus m_χ are displayed in Fig. 5, where $s = \sqrt{\sigma_{\chi\nu}\sigma_{\chi e}}$ for the experimental limit at $n_\sigma = 1.64$ [90% confidence level (CL)] for

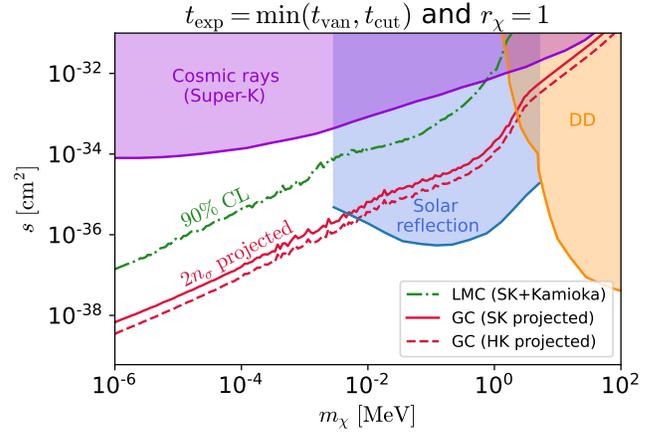


FIG. 5. Sensitivity of BDM searches in $(m_\chi-s)$ plane. $s = \sqrt{\sigma_{\chi\nu}\sigma_{\chi e}}$ for this work on the experimental limits at $n_\sigma = 1.64$ (90% CL) for LMC (green) and projected s for the GC at $n_\sigma = 2$ (red). Current leading limits from cosmic-ray BDM [15], solar reflection [58], and DD [9–12] at $s = \sigma_{\chi e}$ are superimposed.

SN1987a in LMC, and the projected sensitivity at $n_\sigma = 2.0$ for a SN in the GC. In order to compare s with existing constraints based exclusively on $\sigma_{\chi e}$ [9–12, 15, 58], a model-dependent choice relating $\sigma_{\chi\nu}$ and $\sigma_{\chi e}$ has to be made¹. Under a generic description of $r_\chi = \sigma_{\chi\nu}/\sigma_{\chi e}$, the specific case of $r_\chi = 1$ was selected as illustration, with which the resulting bounds are superimposed in Fig. 5. The Super-K constraints are derived from the average background rates [56] and statistical uncertainties. Time stability can be inferred from the absence of anomalous time variations in the solar ν annual modulation analysis [59, 60]. Limits derived with BDM from SN1987a in LMC leads to orders of magnitude improvement over existing bounds for $m_\chi < 2$ keV over a large range of r_χ . For instance, more stringent limits are derived at $m_\chi \sim 10^{-6}$ MeV for $r_\chi > 10^{-6}$. Moreover, a future SN in the GC can improve the sensitivity by a factor of ~ 30 with Super-K, since $s_\chi^{\text{GC}}/s_\chi^{\text{LMC}} \sim \sqrt{N_\chi^{\text{LMC}}/N_\chi^{\text{GC}}} \sim \mathcal{O}(0.03)$, allowing one to probe $s \lesssim 10^{-36}$ cm² for $m_\chi \lesssim 10$ keV. For $m_\chi \lesssim 100$ keV, the sensitivity curves follow $s \propto m_\chi^{1/2}$ simply due to $N_\chi \propto m_\chi^{-1}$ (see Fig. 4).

On the other hand, the sensitivities for $m_\chi \gtrsim 100$ keV weaken considerably due to the reduced BDM that can arrive at the detector within 35 years. Finally, we include an additional projection with Hyper-K for the GC case

¹ Our bounds are considerably better than those reported in Ref. [30, 31], which are not shown in Fig. 5. Also noted is that such a comparison implicitly assumes a cross section that is independent of the center-of-mass energy, which was adopted similarly in pioneer works for cosmic-ray upscattered dark matter scenario [13–15].

(red-dashed curve). The analysis is similar to Super-K, with fiducial mass and background rate scaled up by a factor of 10, which then leads to another improvement of ~ 2 –3 over the Super-K result.

Summary and prospects.—We have examined the scenario of halo DM being boosted by prompt $\text{SN}\nu$, and extracted a wealth of information from its TOF measurements. The BDM events on Earth are characterized by unique timing distributions, which vanish beyond m_χ -dependent end points and are independent of the interaction cross sections, while their peak positions provide information on the SN locations and m_χ .

A new constraint was derived on $s = \sqrt{\sigma_{\chi\nu}\sigma_{\chi e}}$ using Kamiokande and Super-K data on the SN1987a in LMC. Our results probe and exclude new parameter space over a large range of r_χ and in particular improve over the existing cosmic-ray BDM bounds for $m_\chi < 100$ keV by 1–3 orders of magnitude at $\sigma_{\chi e} \sim \sigma_{\chi\nu}$. A future SN in the GC can provide improved sensitivity by another factor of 30–100 with Super-K or Hyper-K. The improvement over other probes [9–12, 15, 30, 31, 58] in the sub-MeV mass range originates from the transient BDM flux arriving in a short duration that can be calibrated by the detection of $\text{SN}\nu$, thereby minimizing the background counts. The constraint and sensitivity of this work were derived by a conservative analysis which stands on the BDM rates not being larger than those of background. A detailed analysis that optimally exploits the m_χ -dependent TOF temporal profile or combines multiple detectors is beyond the scope of this work but will further enhance the sensitivities. Furthermore, most BDM arriving on Earth are within a small solid angle relative to the SN direction for $m_\chi \lesssim \mathcal{O}(\text{MeV})$ (see Fig. 2). Coupled with the good pointing capability for galactic SN [61–64], the angular information can be exploited to greatly reduce the background.

Other effects such as the distortion of $\text{SN}\nu$ spectra, the recently proposed $\text{SN}\nu$ echo [65], and the impact of $\chi - \nu$ interaction on $\text{SN}\nu$ emission have been neglected here. Estimations suggest that $\text{SN}\nu$ spectra be minimally affected for the parameter space examined. These effects may be combined with the TOF profiles of $\text{SN}\nu$ BDM to provide severe constraints on specific phenomenological models relating $\sigma_{\chi e}$ and $\sigma_{\chi\nu}$. With all the rich information, the next galactic SN will offer new insights to the nature of DM. Furthermore, TOF analysis following SN or other transient astrophysical events can be applied in a similar vein to studies of other exotic physics interactions. A broad range of interesting scenario will be explored in our future research.

Code availability.—We provide a Python package `snorer` [66], which can fully reproduce the results in this Letter. The package is available on both GitHub and PyPI. It offers numerous new features, such as including DM spikes, user-specified SN locations in arbitrary distant galaxies, and an implementation of any particle

physics model. See its official page for further details.

We thank Gang Guo, Kin-Wang Ng, Gianluca Petrillo, Yun-Tse Tsai for useful discussions. Y.-H. L. acknowledges the Postdoctoral Scholar Program of the Academia Sinica, Taiwan, during which the major part of this work was done, as well as supports from National Science and Technology Council, the Ministry of Education under Project No. NTU-112L104022, and the National Center for Theoretical Sciences of Taiwan. M.-R. W. acknowledges supports from the National Science and Technology Council, Taiwan under Grant No. 110-2112-M-001-050 and 111-2628-M-001-003-MY4, the Academia Sinica under Project No. AS-CDA-109-M11, and Physics Division, National Center for Theoretical Sciences of Taiwan. H. T.-K. W. acknowledges supports from the National Science and Technology Council, Taiwan under Grant No. 106-2923-M-001-006-MY5. We also acknowledge the computing resources provided by the Academia Sinica Grid-computing Center.

* yehsun@phys.ncku.edu.tw

† mwu@gate.sinica.edu.tw

‡ htwong@phys.sinica.edu.tw

- [1] M. Battaglieri *et al.*, in *U.S. Cosmic Visions: New Ideas in Dark Matter* (2017) [arXiv:1707.04591](https://arxiv.org/abs/1707.04591) [hep-ph].
- [2] R. L. Workman *et al.* (Particle Data Group), [PTEP 2022, 083C01](https://arxiv.org/abs/2022.083C01) (2022).
- [3] M. Aguilar *et al.* (AMS), [Phys. Rev. Lett. 115, 211101](https://arxiv.org/abs/1511.02473) (2015).
- [4] D. S. Akerib *et al.* (LUX), [Phys. Rev. Lett. 118, 021303](https://arxiv.org/abs/1608.07648) (2017), [arXiv:1608.07648](https://arxiv.org/abs/1608.07648) [astro-ph.CO].
- [5] M. Ackermann *et al.* (Fermi-LAT), [Astrophys. J. 840, 43](https://arxiv.org/abs/1704.03910) (2017), [arXiv:1704.03910](https://arxiv.org/abs/1704.03910) [astro-ph.HE].
- [6] D. S. Akerib *et al.* (LUX), [Phys. Rev. Lett. 118, 251302](https://arxiv.org/abs/1705.03380) (2017), [arXiv:1705.03380](https://arxiv.org/abs/1705.03380) [astro-ph.CO].
- [7] E. Aprile *et al.* (XENON), [Phys. Rev. Lett. 121, 111302](https://arxiv.org/abs/1805.12562) (2018), [arXiv:1805.12562](https://arxiv.org/abs/1805.12562) [astro-ph.CO].
- [8] G. Ambrosi *et al.* (DAMPE), [Nature 552, 63](https://arxiv.org/abs/1711.10981) (2017), [arXiv:1711.10981](https://arxiv.org/abs/1711.10981) [astro-ph.HE].
- [9] E. Aprile *et al.* (XENON), [Phys. Rev. Lett. 123, 251801](https://arxiv.org/abs/1907.11485) (2019), [arXiv:1907.11485](https://arxiv.org/abs/1907.11485) [hep-ex].
- [10] E. Aprile *et al.* (XENON), [Phys. Rev. Lett. 123, 241803](https://arxiv.org/abs/1907.12771) (2019), [arXiv:1907.12771](https://arxiv.org/abs/1907.12771) [hep-ex].
- [11] O. Abramoff *et al.* (SENSEI), [Phys. Rev. Lett. 122, 161801](https://arxiv.org/abs/1901.10478) (2019), [arXiv:1901.10478](https://arxiv.org/abs/1901.10478) [hep-ex].
- [12] R. Agnese *et al.* (SuperCDMS), [Phys. Rev. Lett. 121, 051301](https://arxiv.org/abs/1804.10697) (2018), [Erratum: [Phys.Rev.Lett. 122, 069901](https://arxiv.org/abs/1804.10697) (2019)], [arXiv:1804.10697](https://arxiv.org/abs/1804.10697) [hep-ex].
- [13] T. Bringmann and M. Pospelov, [Phys. Rev. Lett. 122, 171801](https://arxiv.org/abs/1810.10543) (2019), [arXiv:1810.10543](https://arxiv.org/abs/1810.10543) [hep-ph].
- [14] Y. Ema, F. Sala, and R. Sato, [Phys. Rev. Lett. 122, 181802](https://arxiv.org/abs/1811.00520) (2019), [arXiv:1811.00520](https://arxiv.org/abs/1811.00520) [hep-ph].
- [15] C. V. Cappiello and J. F. Beacom, [Phys. Rev. D 100, 103011](https://arxiv.org/abs/1906.11283) (2019), [Erratum: [Phys.Rev.D 104, 069901](https://arxiv.org/abs/1906.11283) (2021)], [arXiv:1906.11283](https://arxiv.org/abs/1906.11283) [hep-ph].
- [16] J. B. Dent, B. Dutta, J. L. Newstead, and I. M. Shoemaker, [Phys. Rev. D 101, 116007](https://arxiv.org/abs/1907.03782) (2020), [arXiv:1907.03782](https://arxiv.org/abs/1907.03782) [hep-ph].

- [17] W. Wang, L. Wu, J. M. Yang, H. Zhou, and B. Zhu, *JHEP* **12**, 072 (2020), [Erratum: *JHEP* 02, 052 (2021)], [arXiv:1912.09904 \[hep-ph\]](#).
- [18] Y. Zhang, *PTEP* **2022**, 013B05 (2022), [arXiv:2001.00948 \[hep-ph\]](#).
- [19] G. Guo, Y.-L. S. Tsai, and M.-R. Wu, *JCAP* **10**, 049 (2020), [arXiv:2004.03161 \[astro-ph.HE\]](#).
- [20] S.-F. Ge, J. Liu, Q. Yuan, and N. Zhou, *Phys. Rev. Lett.* **126**, 091804 (2021), [arXiv:2005.09480 \[hep-ph\]](#).
- [21] Q.-H. Cao, R. Ding, and Q.-F. Xiang, *Chin. Phys. C* **45**, 045002 (2021), [arXiv:2006.12767 \[hep-ph\]](#).
- [22] Y. Jho, J.-C. Park, S. C. Park, and P.-Y. Tseng, *Phys. Lett. B* **811**, 135863 (2020), [arXiv:2006.13910 \[hep-ph\]](#).
- [23] W. Cho, K.-Y. Choi, and S. M. Yoo, *Phys. Rev. D* **102**, 095010 (2020), [arXiv:2007.04555 \[hep-ph\]](#).
- [24] Z.-H. Lei, J. Tang, and B.-L. Zhang, (2020), [arXiv:2008.07116 \[hep-ph\]](#).
- [25] G. Guo, Y.-L. S. Tsai, M.-R. Wu, and Q. Yuan, *Phys. Rev. D* **102**, 103004 (2020), [arXiv:2008.12137 \[astro-ph.HE\]](#).
- [26] C. Xia, Y.-H. Xu, and Y.-F. Zhou, *Nucl. Phys. B* **969**, 115470 (2021), [arXiv:2009.00353 \[hep-ph\]](#).
- [27] J. B. Dent, B. Dutta, J. L. Newstead, I. M. Shoemaker, and N. T. Arellano, *Phys. Rev. D* **103**, 095015 (2021), [arXiv:2010.09749 \[hep-ph\]](#).
- [28] Y. Ema, F. Sala, and R. Sato, *SciPost Phys.* **10**, 072 (2021), [arXiv:2011.01939 \[hep-ph\]](#).
- [29] V. V. Flambaum, L. Su, L. Wu, and B. Zhu, (2020), [arXiv:2012.09751 \[hep-ph\]](#).
- [30] Y. Jho, J.-C. Park, S. C. Park, and P.-Y. Tseng, (2021), [arXiv:2101.11262 \[hep-ph\]](#).
- [31] A. Das and M. Sen, *Phys. Rev. D* **104**, 075029 (2021), [arXiv:2104.00027 \[hep-ph\]](#).
- [32] N. F. Bell, J. B. Dent, B. Dutta, S. Ghosh, J. Kumar, J. L. Newstead, and I. M. Shoemaker, *Phys. Rev. D* **104**, 076020 (2021), [arXiv:2108.00583 \[hep-ph\]](#).
- [33] W. Chao, T. Li, and J. Liao, (2021), [arXiv:2108.05608 \[hep-ph\]](#).
- [34] D. Ghosh, A. Guha, and D. Sachdeva, *Phys. Rev. D* **105**, 103029 (2022), [arXiv:2110.00025 \[hep-ph\]](#).
- [35] J.-C. Feng, X.-W. Kang, C.-T. Lu, Y.-L. S. Tsai, and F.-S. Zhang, *JHEP* **04**, 080 (2022), [arXiv:2110.08863 \[hep-ph\]](#).
- [36] W. Wang, L. Wu, W.-N. Yang, and B. Zhu, (2021), [arXiv:2111.04000 \[hep-ph\]](#).
- [37] C. Xia, Y.-H. Xu, and Y.-F. Zhou, *JCAP* **02**, 028 (2022), [arXiv:2111.05559 \[hep-ph\]](#).
- [38] J.-W. Wang, A. Granelli, and P. Ullio, *Phys. Rev. Lett.* **128**, 221104 (2022), [arXiv:2111.13644 \[astro-ph.HE\]](#).
- [39] X. Cui *et al.* (PandaX-II), *Phys. Rev. Lett.* **128**, 171801 (2022), [arXiv:2112.08957 \[hep-ex\]](#).
- [40] R. Xu *et al.* (CDEX), (2022), [arXiv:2201.01704 \[hep-ex\]](#).
- [41] A. Granelli, P. Ullio, and J.-W. Wang, (2022), [arXiv:2202.07598 \[astro-ph.HE\]](#).
- [42] W. D. Arnett and J. L. Rosner, *Phys. Rev. Lett.* **58**, 1906 (1987).
- [43] E. W. Kolb, A. J. Stebbins, and M. S. Turner, *Phys. Rev. D* **35**, 3598 (1987).
- [44] B. Holdom, *Phys. Lett. B* **166**, 196 (1986).
- [45] C. A. Argüelles, A. Kheirandish, and A. C. Vincent, *Phys. Rev. Lett.* **119**, 201801 (2017), [arXiv:1703.00451 \[hep-ph\]](#).
- [46] J. H. Chang, R. Essig, and S. D. McDermott, *JHEP* **09**, 051 (2018), [arXiv:1803.00993 \[hep-ph\]](#).
- [47] K. Murase and I. M. Shoemaker, *Phys. Rev. Lett.* **123**, 241102 (2019), [arXiv:1903.08607 \[hep-ph\]](#).
- [48] G.-L. Lin and Y.-H. Lin, *Phys. Rev. D* **104**, 063021 (2021), [arXiv:2102.11151 \[hep-ph\]](#).
- [49] D. Croon, G. Elor, R. K. Leane, and S. D. McDermott, *JHEP* **01**, 107 (2021), [arXiv:2006.13942 \[hep-ph\]](#).
- [50] M. Escudero, D. Hooper, G. Krnjaic, and M. Pierre, *JHEP* **03**, 071 (2019), [arXiv:1901.02010 \[hep-ph\]](#).
- [51] P. Foldenauer, *Phys. Rev. D* **99**, 035007 (2019), [arXiv:1808.03647 \[hep-ph\]](#).
- [52] H. Duan, G. M. Fuller, J. Carlson, and Y.-Z. Qian, *Phys. Rev. D* **74**, 105014 (2006), [arXiv:astro-ph/0606616](#).
- [53] G. Bertone, D. Hooper, and J. Silk, *Phys. Rept.* **405**, 279 (2005), [arXiv:hep-ph/0404175](#).
- [54] D. Erkal *et al.*, *Mon. Not. Roy. Astro.* **487**, 2685 (2019).
- [55] K. Hirata *et al.*, *Phys. Rev. Lett.* **58**, 1490 (1987).
- [56] K. Abe *et al.* (Super-Kamiokande), *Phys. Rev. D* **94**, 052010 (2016), [arXiv:1606.07538 \[hep-ex\]](#).
- [57] G. Battistoni, A. Ferrari, T. Montaruli, and P. R. Sala, *Astropart. Phys.* **23**, 526 (2005).
- [58] H. An, M. Pospelov, J. Pradler, and A. Ritz, *Phys. Rev. Lett.* **120**, 141801 (2018), [Erratum: *Phys.Rev.Lett.* 121, 259903 (2018)], [arXiv:1708.03642 \[hep-ph\]](#).
- [59] J. Yoo *et al.* (Super-Kamiokande), *Phys. Rev. D* **68**, 092002 (2003), [arXiv:hep-ex/0307070](#).
- [60] J. P. Cravens *et al.* (Super-Kamiokande), *Phys. Rev. D* **78**, 032002 (2008), [arXiv:0803.4312 \[hep-ex\]](#).
- [61] J. F. Beacom and P. Vogel, *Phys. Rev. D* **60**, 033007 (1999), [arXiv:astro-ph/9811350](#).
- [62] R. Tomas, D. Semikoz, G. G. Raffelt, M. Kachelriess, and A. S. Dighe, *Phys. Rev. D* **68**, 093013 (2003), [arXiv:hep-ph/0307050](#).
- [63] K. Abe *et al.* (Super-Kamiokande), *Astropart. Phys.* **81**, 39 (2016), [arXiv:1601.04778 \[astro-ph.HE\]](#).
- [64] N. B. Linzer and K. Scholberg, *Phys. Rev. D* **100**, 103005 (2019), [arXiv:1909.03151 \[astro-ph.IM\]](#).
- [65] J. A. Carpio, A. Kheirandish, and K. Murase, (2022), [arXiv:2204.09650 \[hep-ph\]](#).
- [66] GitHub: <https://github.com/yenhunlin/snorer> and PyPI: <https://pypi.org/project/snorer/>.